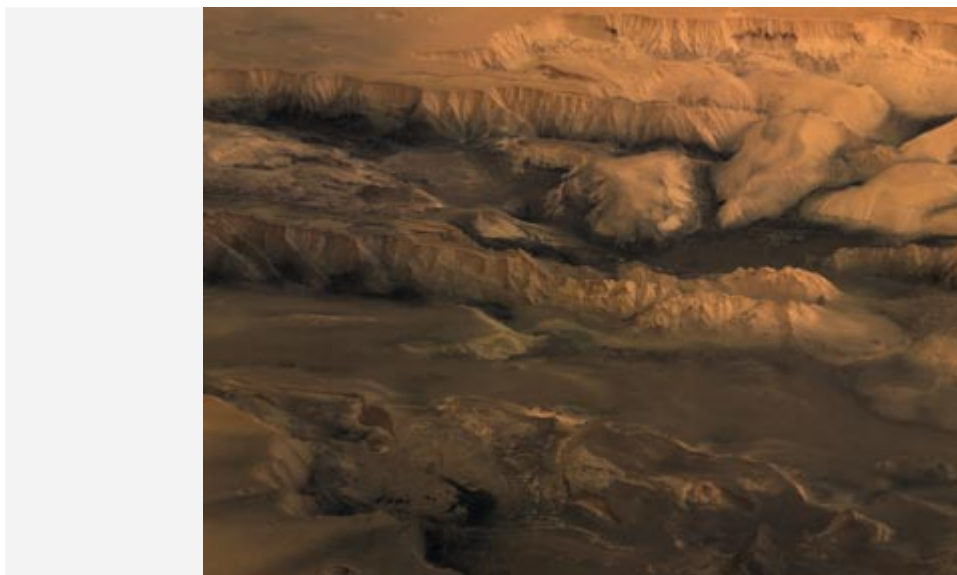


News Archive until 2006

Melas Chasma, Candor Chasma and Ophir Chasma: The central part of Valles Marineris

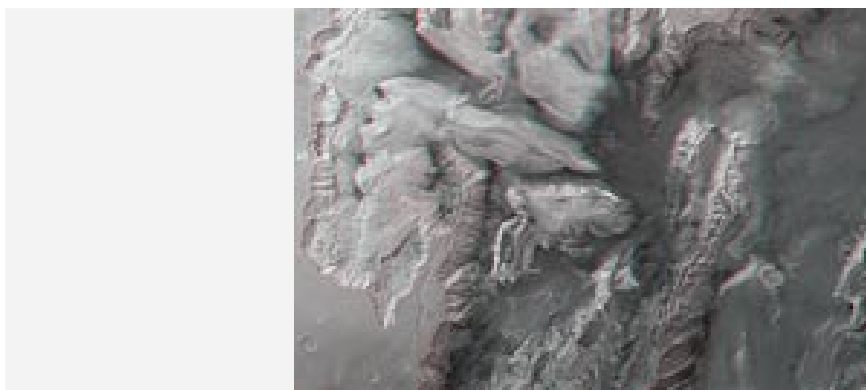
28 August 2006



Melas, Candor and Ophir Chasma, colour perspective view

These images, taken by the High Resolution Stereo Camera (HRSC) onboard ESA's Mars Express spacecraft, show the central part of the 4000-kilometre long Valles Marineris canyon on Mars. The HRSC is operated by and was developed by the German Aerospace Center (DLR).

The HRSC obtained these images during orbits 334 and 360 with a resolution of approximately 21 metres per pixel for the earlier orbit and 30 metres per pixel for the latter.



Melas, Candor and Ophir Chasma, 3D image

The scenes show an area of approximately 300 by 600 kilometres and are taken from an image mosaic that was created from the two orbit sequences. The image above is located between 3° to 13° South, and 284° to 289° East.

Valles Marineris was named after the US Mariner 9 probe, the first spacecraft to image this enormous feature in 1971. Here, the huge canyon which runs east to west is at its widest in the north-south direction.

It remains unclear how this gigantic geological feature, unparalleled in the Solar System, was formed. Tensions in the upper crust of Mars possibly led to cracking of the highlands. Subsequently, blocks of the crust slid down between these tectonic fractures.

The fracturing of Valles Marineris could have occurred thousands of millions of years ago, when the Tharsis bulge (west of Valles Marineris) began to form as the result of volcanic activity and subsequently grew to the dimensions of greater than a thousand kilometres in diameter and more than ten kilometres high. On Earth, such a tectonic process is called 'rifting', presently occurring on a smaller scale in the Kenya rift in eastern Africa.



Melas, Candor and Ophir Chasma in colour



Melas, Candor and Ophir Chasmas in black and white

The collapse of large parts of the highland is an alternative explanation. For instance, extensive amounts of water ice could have been stored beneath the surface and were then melted as a result of thermal activity, most likely the nearby volcanic Tharsis province.

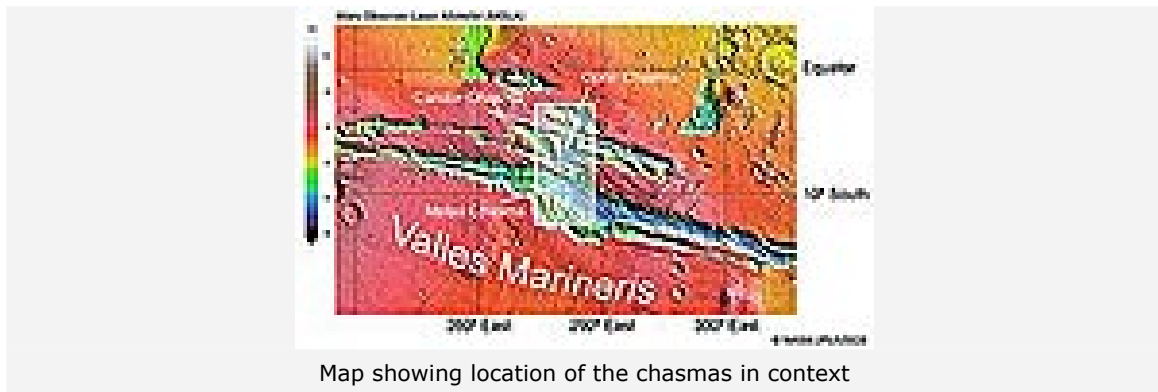
The water could have travelled towards the northern lowlands, leaving cavities beneath the surface where the ice once existed. The roofs could no longer sustain the load of the overlying rocks, so the area collapsed.

Regardless of how Valles Marineris might have formed, it is clear that once the depressions were formed and the surface was topographically structured, heavy erosion then began shaping the landscape.

Two distinct landforms can be distinguished. On one hand, we see sheer cliffs with prominent edges and ridges. These are erosion features that are typical in arid mountain zones on Earth.

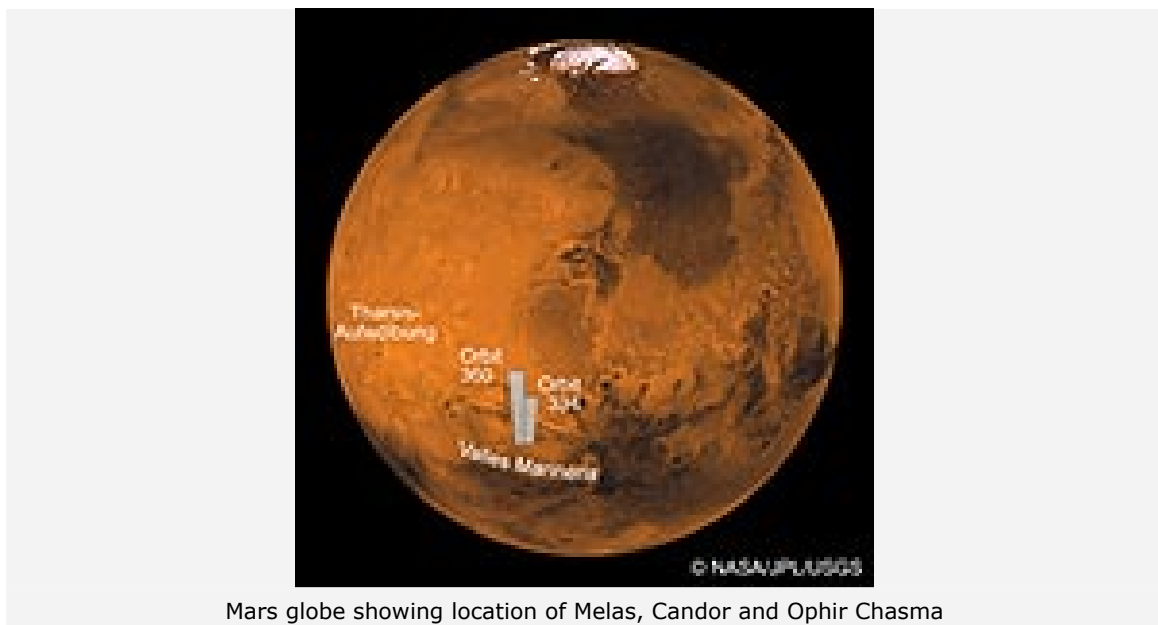
Today, the surface of Mars is bone dry, so wind and gravity are the dominant processes that shape the landscape (this might have been much different in the geological past of the planet when Valles Marineris possibly had flowing water or glaciers winding down its slopes).

In contrast, some gigantic 'hills' (indeed, between 1000 and 2000 metres high) located on the floors of the valleys have a smoother topography and a more sinuous outline. So far, scientists have no definitive explanation for why these different landforms exist.



Below the northern scarp, there are several landslides, where material was transported over a distance of up to 70 kilometres. Also seen in the image there are several structures suggesting flow of material in the past. Therefore, material could have been deposited in the valleys, making the present floor look heterogeneous.

In the centre of the larger images above, there are surface features that appear similar to ice flows. These were previously identified in pictures from the US Viking probes of the 1970s; their origin remains a mystery.



The colour images were processed using the HRSC nadir (vertical view) and three colour channels. The perspective views were calculated from the digital terrain model derived from the stereo channels.

The 3D anaglyph image was created from the nadir channel and one of the stereo channels. Stereoscopic glasses are needed to view the 3D image. Image resolution has been decreased for use on the internet.

HRSC on the ESA Mars Express mission

The High Resolution Stereo Camera (HRSC) on the European Space Agency (ESA) Mars Express mission was developed by the German Aerospace Center (DLR) and built in cooperation with industrial partners (EADS Astrium, Lewicki Microelectronic GmbH and Jena-Optronik GmbH). The HRSC experiment on Mars Express is led by the Principal Investigator (PI) Prof. Dr Gerhard Neukum. The science team of the experiment consists of 45 Co-Investigators from 32 institutions and 10 nations.

Related Contacts

Prof.Dr. Ralf Jaumann

German Aerospace Center

Institute of Planetary Research, Planetary Geology

Tel: +49 30 67055-400

Fax: +49 30 67055-402
E-Mail: Ralf.Jaumann@dlr.de

Marco Trovatiello

German Aerospace Center (DLR)
Corporate Communications
Tel: +49 2203 601-2116
Fax: +49 2203 601-3249
E-Mail: marco.trovatiello@dlr.de

Prof. Dr. Gerhard Neukum

FU Berlin, Institut of Geological Sciences
Tel: +49 30 838-70579/-70575
E-Mail: gneukum@zedat.fu-berlin.de

Contact details for image and video enquiries as well as information regarding DLR's terms of use can be found on the DLR portal imprint.